

Open Research Online

The Open University's repository of research publications
and other research outputs

Herschel-SPIRE spectroscopy of the DR21 molecular cloud core

Journal Item

How to cite:

White, G. J.; Abergel, A.; Spencer, L.; Schneider, N.; Naylor, D. A.; Anderson, L. D.; Joblin, C.; Ade, P.; André, P.; Arab, H.; Baluteau, J.-P.; Bernard, J.-P.; Blagrove, K.; Bontemps, S.; Boulanger, F.; Cohen, M.; Compiegne, M.; Cox, P.; Dartois, E.; Davis, G.; Emery, R.; Fulton, T.; Gom, B.; Griffin, M.; Gry, C.; Habart, E.; Huang, M.; Jones, S.; Kirk, J. M.; Lagache, G.; Leeks, S.; Lim, T.; Madden, S.; Makiwa, G.; Martin, P.; Miville-Deschênes, M.-A.; Molinari, S.; Moseley, H.; Motte, F.; Okumura, K.; Pinheiro Gonçalves, D.; Polehampton, E.; Rodet, T.; Rodón, J. A.; Russeil, D.; Saraceno, P.; Sidher, S.; Swinyard, B. M.; Ward-Thompson, D. and Zavagno, A. (2010). *Herschel*-SPIRE spectroscopy of the DR21 molecular cloud core. *Astronomy & Astrophysics*, 518, article no. L114.

For guidance on citations see [FAQs](#).

© 2010 ESO

Version: Version of Record

Link(s) to article on publisher's website:

<http://dx.doi.org/doi:10.1051/0004-6361/201014622>

Copyright and Moral Rights for the articles on this site are retained by the individual authors and/or other copyright owners. For more information on Open Research Online's data [policy](#) on reuse of materials please consult the policies page.

oro.open.ac.uk

LETTER TO THE EDITOR

***Herschel*-SPIRE spectroscopy of the DR21 molecular cloud core[★]**

G. J. White^{1,2}, A. Abergel⁴, L. Spencer³, N. Schneider⁵, D. A. Naylor³, L.D. Anderson⁸, C. Joblin^{17,18}, P. Ade⁷, P. André⁵, H. Arab⁴, J.-P. Baluteau⁸, J.-P. Bernard^{17,18}, K. Blagrove⁹, S. Bontemps^{5,6}, F. Boulanger⁴, M. Cohen¹⁰, M. Compiegne⁹, P. Cox¹¹, E. Dartois⁴, G. Davis¹², R. Emery², T. Fulton⁴, B. Gom⁴, M. Griffin⁷, C. Gry⁸, E. Habart⁴, M. Huang¹⁴, S. Jones³, J. M. Kirk⁷, G. Lagache⁴, S. Leeks², T. Lim², S. Madden⁵, G. Makiwa³, P. Martin⁹, M.-A. Miville-Deschênes⁴, S. Molinari¹⁴, H. Moseley¹⁵, F. Motte⁵, K. Okumura⁵, D. Pinheiro Gonçalves⁹, E. Polehampton^{2,4}, T. Rodet¹⁶, J. A. Rodón⁸, D. Russeil⁸, P. Saraceno¹⁴, S. Sidher², B. M. Swinyard², D. Ward-Thompson⁷, A. Zavagno⁸

(Affiliations are available in the online edition)

Received 31 March 2010 / Accepted 11 May 2010

ABSTRACT

We present far-infrared spectra and maps of the DR21 molecular cloud core between 196 and 671 μm , using the *Herschel*-SPIRE spectrometer. Nineteen molecular lines originating from CO, ^{13}CO , HCO^+ and H_2O , plus lines of [N II] and [CI] were recorded, including several transitions not previously detected. The CO lines are excited in warm gas with $T_{\text{kin}} \sim 125$ K and $n_{\text{H}_2} \sim 7 \times 10^4 \text{ cm}^{-3}$, CO column density $N(\text{CO}) \sim 3.5 \times 10^{18} \text{ cm}^{-2}$ and a filling factor of $\sim 12\%$, and appear to trace gas associated with an outflow. The rotational temperature analysis incorporating observations from ground-based telescopes reveals an additional lower excitation CO component which has a temperature ~ 78 K and $N(\text{CO}) \sim 4.5 \times 10^{21} \text{ cm}^{-2}$.

Key words. ISM: lines and bands – infrared: ISM – ISM: molecules

1. Introduction

We report observations of the far-IR spectrum of the DR21 molecular cloud core obtained with the *Herschel* satellite between 196 and 671 μm . The DR21 HII-region/molecular cloud is part of the Cygnus X complex of molecular clouds located at a distance of 1.7 kpc (Schneider et al. 2006). This region has been subject to numerous studies at different wavelengths (Richardson et al. 1988; Wilson & Mauersberger 1990; Liechti & Walmsley 1997; Schneider et al. 2006, 2010; Jakob et al. 2007). The main DR21 cloud core has a mass of $\sim 20\,000 M_{\odot}$ (Richardson et al. 1989), and contains one of the most energetic star formation outflows detected, with an outflow mass of $\sim 3000 M_{\odot}$ (Garden et al. 1991; Cruz-González et al. 2010).

2. SPIRE observations

2.1. Spectra

We present science demonstration phase (SDP) observations obtained with ESA's *Herschel* Space Observatory (Pilbratt et al. 2010), using the Spectral and Photometric Imaging REceiver (SPIRE – Griffin et al. 2010). The calibration and characteristics of SPIRE have been described by Swinyard et al. (2010). SPIRE was operated as an imaging Fourier-transform spectrometer (FTS) in the high resolution mode ($\lambda/\Delta\lambda = 1000$ ($\sim 300 \text{ km s}^{-1}$

at 250 μm) sampling across an approximately circular field of view with an unvignetted diameter of 2.6'. This means that the line profiles are unresolved.

The sky footprint is formed by two detector arrays: the 19 pixel SLW array (671–303 μm) and the 37 pixel SSW array (313–194 μm), with beam widths varying from 17'' at 194 μm to 42'' at 671 μm , with uncertainties of ± 7 –10% (Griffin et al. 2010). The integration time was 1065 s, summed from two separate observations. The current best estimates of the absolute uncertainties for the FTS detectors are 10–20% for the SSW detectors, and $\sim 30\%$ for the SLW detectors (Swinyard et al. 2010).

The unapodised FTS spectra provide the highest spectral resolution, with a classical instrumental sinc function line shape. A spectral line fitting routine was developed for extracting line parameters (Jones et al. 2009). This fits a continuum (either a low order polynomial or a blackbody variant) using the Levenberg-Marquardt least squares method. The fitting procedure weights the spectral intensity at a given frequency of an averaged spectrum by the statistical uncertainty at that frequency, returning line centers, intensities, line widths and their associated fit errors.

2.2. Maps

The SPIRE observations sparsely sample the field of view, although there are calibration uncertainties for the outer ring of detectors at the edges of both arrays that are not yet fully characterised. To provide a first look at the relative distributions in the various species, we have interpolated the fluxes of individual

[★] *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

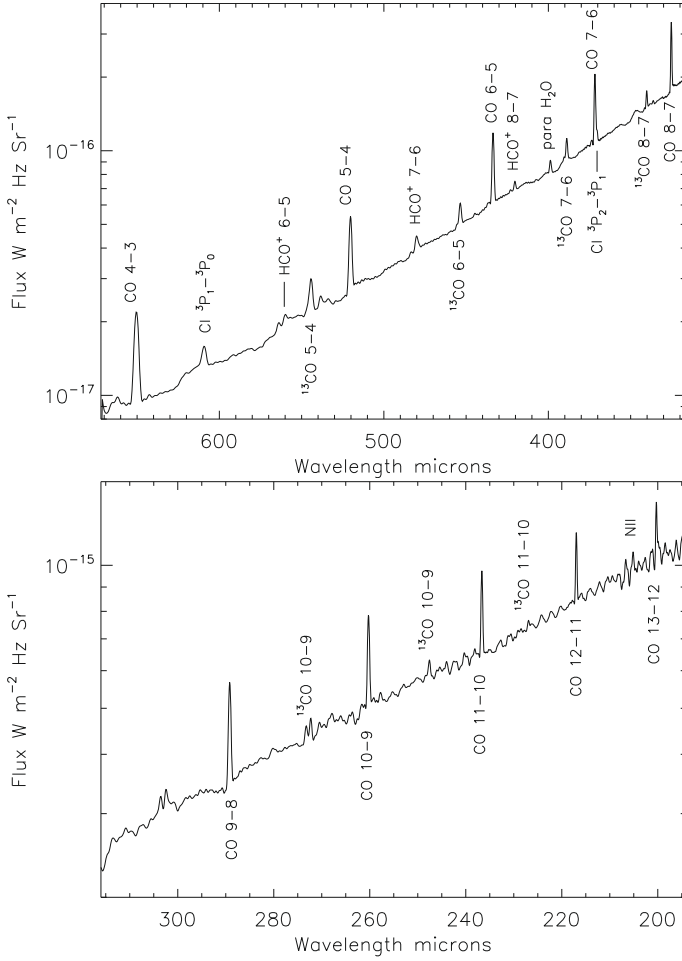


Fig. 1. (Upper picture) Spectrum of DR21 with the SLW detector array, with beam sizes varying from 42.0'' at 671.1 μm to 37.3'' at 316.5 μm , and (lower picture) with the SSW detector array, with beam sizes varying from 18.7'' at 281.7 μm to 16.8'' at 194.2 μm .

lines, although this spatial information is not used in subsequent line modeling.

The maps of selected species are shown in Fig. 2. The CO lines in both detector arrays show a prominent central peak, with extensions to the east and west along the well known outflow. This has been assumed to be associated with outflowing gas with ($T_{\text{ex}} \sim 2000$ K and $H(\text{H}_2) \sim 1 \times 10^{19} \text{ cm}^{-2}$ from Garden et al. 1991). However, as will be seen in the high resolution JCMT observations (Fig. 3), the emission traced in the SPIRE maps is also clearly visible in the relatively low excitation CO $J = 3-2$ data, suggesting that there may be a mixture of low and high excitation gas present. This is confirmed in Fig. 3, where similar extensions of the ambient gas are present in the JCMT CO $J = 3-2$ map, and that of Schneider et al. (2010). This is not unexpected, as this outflow appears to have a very large mass of several thousand M_{\odot} , and presumably the high velocity gas phase overlaps (or may co-exist with) ambient material. The SPIRE maps also show that the $^3\text{P}_1-^3\text{P}_0$ atomic carbon line has a similar spatial distribution to that of CO. By contrast, the H_2O and [N II] lines appear to be more compact and centred close to the DR21 cloud core, although the [N II] distribution is elongated to the east – observations with higher signal to noise and better sampling are needed for more detailed comparison.

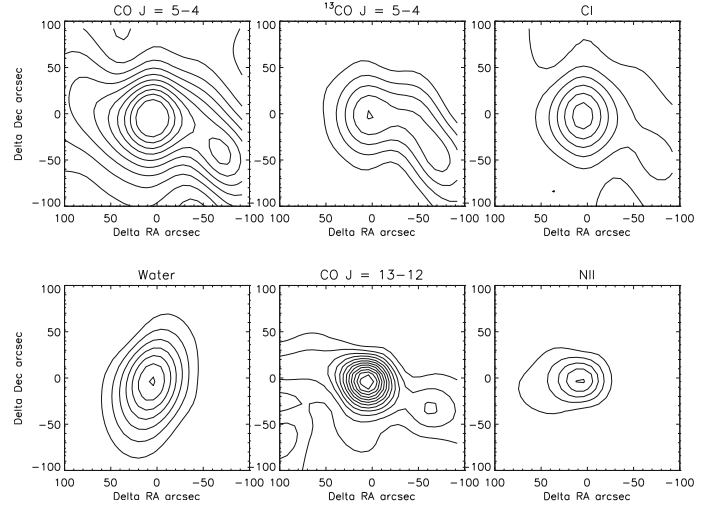


Fig. 2. The SPIRE maps of selected species toward DR21. Since the data points do not fully sample the mapped area, this map was constructed by interpolating between the data points. The (0,0) position is located at RA (2000) = 20^h39^m01^s.18, Dec (2000) = +42° 19' 43".66.

Table 1. Fluxes measured in the central pixel.

Species	Transition	Wave μm	Int intensity $\text{W m}^{-2} \text{sr}^{-1}$	Intens error $\text{W m}^{-2} \text{sr}^{-1}$
CO	$J = 4-3$	650.1	2.85 (–8)	6.93 (–10)
CI	$^3\text{P}_1-^3\text{P}_0$	609.0	4.86 (–9)	9.96 (–10)
HCO ⁺	$J = 6-5$	560.5	3.99 (–9)	4.29 (–10)
¹³ CO	$J = 5-4$	544.1	1.66 (–8)	5.04 (–10)
CO	$J = 5-4$	520.3	6.81 (–8)	3.39 (–10)
HCO ⁺	$J = 7-6$	480.3	1.02 (–8)	1.35 (–9)
¹³ CO	$J = 6-5$	453.5	2.44 (–8)	3.21 (–9)
CO	$J = 6-5$	433.5	1.15 (–7)	1.47 (–8)
HCO ⁺	$J = 8-7$	420.3	1.32 (–8)	2.10 (–9)
H ₂ O	$2_{11}-2_{02}$	398.6	2.33 (–8)	3.03 (–9)
¹³ CO	$J = 7-6$	388.7	3.66 (–8)	5.88 (–9)
CO	$J = 7-6$	371.6	2.14 (–7)	1.29 (–9)
CI	$^3\text{P}_2-^3\text{P}_1$	370.5	3.03 (–8)	1.26 (–9)
¹³ CO	$J = 8-7$	340.1	6.79 (–8)	1.80 (–8)
CO	$J = 8-7$	325.2	3.15 (–7)	4.56 (–8)
CO	$J = 9-8$	289.1	4.89 (–7)	4.23 (–9)
CO	$J = 10-9$	260.2	5.94 (–7)	1.01 (–8)
CO	$J = 11-10$	236.6	7.26 (–7)	5.46 (–9)
CO	$J = 12-11$	216.9	7.44 (–7)	6.72 (–9)
NII	$^3\text{P}_1-^3\text{P}_0$	205.2	1.45 (–7)	4.71 (–8)
CO	$J = 13-12$	200.3	6.90 (–7)	3.96 (–8)

2.3. JCMT CO $J = 3-2$ observations

CO $J = 3-2$ JCMT archival data (programme M07AU01) with a 15'' beam and spectral of 0.05 km s^{-1} are shown in Fig. 3, from a 4.5 h integration using the HARP array receiver. The area covered by the SPIRE footprint (Fig. 2) is shown as a white square. The JCMT observations clearly trace the outflow which runs from the NE-SW from DR21 from the centre of the white box. The JCMT map also reveals a prominent north-south ridge that includes CO peaks associated with the well-studied sources DR21(OH) and DR21-FIR1. Around the DR21 core, a bipolar structure close to the systemic velocity is coincident with distribution of high velocity gas and shocked H_2 (Garden et al. 1991).

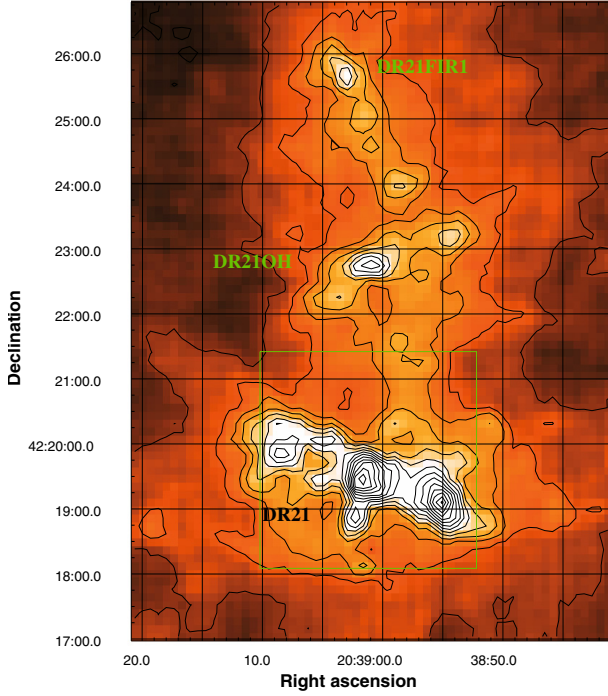


Fig. 3. CO $J = 3-2$ integrated emission map of DR21 between velocities of -40 – ± 40 km s $^{-1}$. The *lowest contour* is drawn at 1 K km s $^{-1}$ (5σ as observed in a 61 kHz (~ 0.05 km s $^{-1}$) spectral channel), and the contour intervals are incremented in steps of 50 K km s $^{-1}$. The peak antenna temperature and velocity integrated temperatures at the DR21 peak are 39.2 K and 726 K km s $^{-1}$ respectively. The temperature scale is calibrated in antenna temperature T_a^* units, which is related to the main beam brightness temperature T_{mb} by $T_{mb} = T_a^* / \eta_{mb}$, where the main beam efficiency, $\eta_{mb} = 0.63$. Further details of HARP calibration are given in Buckle et al. (2010).

3. Modeling the CO lines

The most extensive modeling of the CO toward DR21 is by Richardson et al. (1986, 1988), Wilson et al. (1990), Schneider et al. (2006) and Jakob et al. (2007). Richardson et al. (1986, 1988) presented a multiphase model with gas densities spanning the range 10^3 – 10^6 cm $^{-3}$, and gas temperatures in the low temperature component $\gtrsim 30$ K. Jakob et al. (2007) confirmed this using KOSMA and ISO observations, finding and additional warm phase component with $T_{kin} \sim 80$ – 150 K and clump density $n_{H_2} \sim 10^4$ – 10^6 cm $^{-3}$.

We initially constructed a rotational temperature diagram for the SPIRE CO and ^{13}CO lines. These were augmented with the JCMT CO line from Sect. 2.3, plus IRAM CO $J = 2-1$ observations (Schneider et al. 2010), with suitable beam size corrections. The rotational temperature diagram is shown in Fig. 4.

Both species show evidence for two gas components, a lower temperature phase with a rotational temperature $T_{rot} = 78$ K and total CO column density $N(\text{CO}) \sim 4.5 \times 10^{21}$ cm $^{-2}$, in addition to a higher temperature component with $T_{rot} = 185$ K and $N(\text{CO}) \sim 9.7 \times 10^{17}$ cm $^{-2}$. The ^{13}CO lines are more limited and noisy, with the SPIRE lines indicating an intermediate temperature phase having $T_{rot} = 109$ K and $N(^{13}\text{CO}) \sim 8.8 \times 10^{16}$ cm $^{-2}$. The data for ^{13}CO also show evidence for a low temperature component, although this relies on comparison with low frequency ground based data (JCMT, IRAM) obtained with different beam sizes. Such a result is expected, since the observations probe deeper into the PDR of each clump in ^{13}CO than in CO.

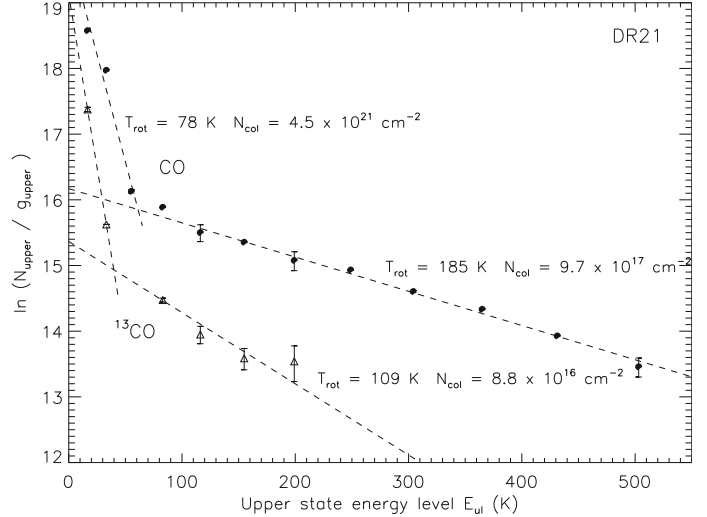


Fig. 4. CO and ^{13}CO line rotational temperature diagram.

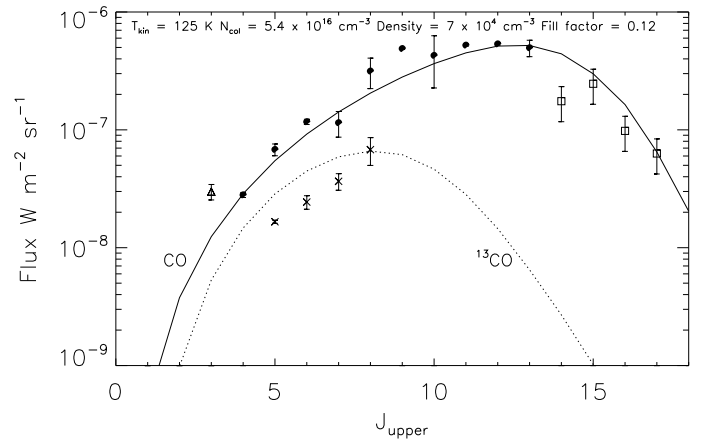


Fig. 5. LVG analysis of the CO and ^{13}CO lines toward DR21. The SPIRE data points, ($J = 4-3$ to $13-12$) are shown as filled circles, and the ISO data points (Jakob et al. 2007) as squares, along with 30% absolute error bars. The CO $J = 3-2$ line was obtained from the JCMT data, convolved to a resolution of $17''$ to match the beamsizes of the nearby SPIRE lines. The filled circles are the SPIRE data, open squares are ISO data from Jakob et al. (2007), open triangle is determined from the JCMT data in this paper, and the crosses are the SPIRE ^{13}CO data. The top solid curve shows the fit to the predicted SPIRE CO line intensities, and the lower dashed one is for ^{13}CO with an assumed CO/ ^{13}CO abundance of 67, an estimated filling factor of 12%. As might be expected, the fit for a single temperature lies intermediate between the two T_{rot} values inferred from Fig. 4.

There are several problems with the rotational temperature approach, including wavelength dependent beam size corrections, opacity and calibration errors. These uncertainties can however be mitigated by i) taking ratios of the various CO line intensities on a single detector and using these to constrain the excitation conditions through our LVG modeling, and ii) using observations from the central pixel where the SSW and SLW beams are coincident and the calibration is well determined. This approach particularly mitigates against the beam size and calibration errors, since only flux ratios are being used to estimate the excitation conditions.

The model fit was made to the CO and ^{13}CO lines using the off-line version of the RADEX LVG code (Van der Tak et al. 2007). The line ratios observed on the same detectors (hence the beam sizes are similar) were used to restrict the likely

excitation conditions. It proved difficult to find an unique single temperature model that simultaneously predicted the relative intensities of both isotopologues. However, the SPIRE data can be approximately reproduced by a single phase moderate temperature gas with $T_{\text{kin}} \sim 125$ K, volume density $\sim 7 \times 10^4 \text{ cm}^{-3}$, with $N(\text{CO}) \sim 3.5 \times 10^{18} \text{ cm}^{-2}$, filling factor $\sim 12\%$, and a $[\text{C}]/[^{13}\text{C}]$ ratio of 65. This model does however slightly overpredict the low- J ^{13}CO 4–3 – 6–5 line intensities, compared to the $J = 7–6$ line. Changing the temperature and density from these conditions considerably worsened the high- J CO line fits, although a more complex multiphase model, with appropriate (and uncertain) beam size corrections would improve the fit of the low J -lines. We have not attempted to fit to a PDR-model, as the data and calibration quality need to be improved if tests between models are to be made, and that this is beyond this first look paper.

4. Modeling the H_2O and $[\text{N II}]$ lines

An objective of this study was to detect the $[\text{N II}]$ 205 μm line, and to compare it with the $[\text{C II}]_{157}$ line which has $n_{\text{crit}} = 46 \text{ cm}^{-3}$, $T_e = 8000$ K. This has a nearly identical critical density for excitation in ionised regions. Their line ratio is directly related to the N^+/C^+ abundance ratio, and this ratio traces the fraction of the observed $[\text{C II}]$ emission that arises from ionized regions (Oberst et al. 2006). Taking the SPIRE upper limit of $7.5 \times 10^{-8} \text{ W m}^{-2} \text{ sr}^{-1}$ with Jakob et al. (2007), the ratio of the 122/205 μm lines is ≥ 1.9 , which is only adequate to constrain the ionised gas density to be $\geq \sim 30 \text{ cm}^{-3}$. The $[\text{C II}]/[\text{N II}]_{205}$ ratio using the Jakob et al. (2007) tabulation is ≥ 5.6 . Given current uncertainties and lack of an $[\text{N II}]_{122}$ flux, it is necessary to await improved data. We note that the $[\text{N II}]$ extension to the east (see Fig. 2) coincides with a hole in the excited H_2 emission image (Cruz-González et al. 2010), which may indicate there is a cavity of ionised gas. However clarification will require future observations with better sampling.

In Fig. 6 we show a section of the spectrum with the 398.6 μm para- H_2O line, and the HCO^+ $J = 6–5$, 7–6 and 8–7 lines. Putting the SPIRE sensitivity into perspective, Jakob et al. (2007) report that the integrated $\text{CI } ^3\text{P}_1-^3\text{P}_0$ intensity measured from the KOSMA telescope with an $80''$ beam is 46.6 K km s^{-1} , and main beam brightness temperature ~ 25 K. By comparison the same line observed with SPIRE has a peak S/N ratio of ≥ 14 as seen in a single SPIRE channel. We also used RADEX to compute an LVG solution for the $2_{11}-2_{02}$ para- H_2O line at 398.5 μm . Assuming similar excitation to that from the CO solution, for an abundance $X[\text{H}_2\text{O}] = \sim 4 \times 10^{-8}$ and line width of 40 km s^{-1} (Hjalmarson et al. 2003), we predict that the SPIRE flux should be $2.4 \times 10^{-8} \text{ W m}^{-2} \text{ sr}^{-1}$, which agrees with the measured value of $2.33 \pm 0.3 \times 10^{-8} \text{ W m}^{-2} \text{ sr}^{-1}$.

5. Conclusions

We have presented the SPIRE spectrum of a star-forming molecular core, DR21, showing for the first time the complete CO and ^{13}CO band head from $J = 4–3$ to $13–12$, along with their maps at far-infrared wavelengths. A rotational temperature analysis shows two gas phases with $T_{\text{rot}} \sim 80$ K and CO column density $\sim 4.5 \times 10^{18} \text{ cm}^{-2}$, and $T_{\text{rot}} = 185$ K and $N(\text{CO}) \sim 10^{18} \text{ cm}^{-2}$

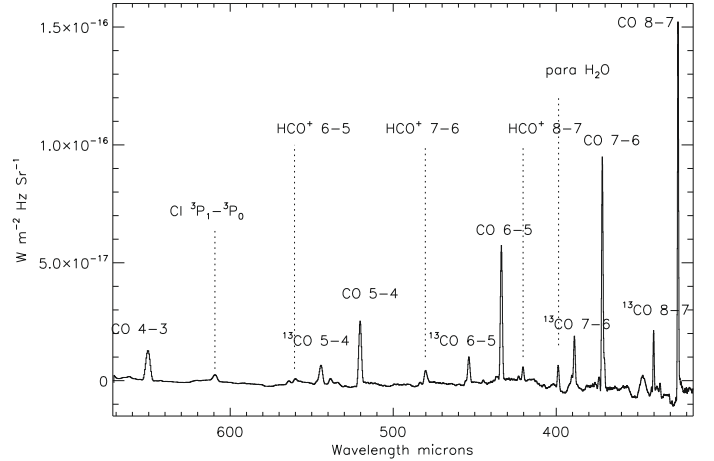


Fig. 6. Continuum subtracted spectrum showing the 398.6 μm para- H_2O , and HCO^+ $J = 7–6$ and $J = 8–7$ lines. Some instrumental features remain – notably a broad bump close to the CH^+ line at 347.9 μm . Although it is possible to flatten this spectrum further, we chose only to fit a third order polynomial to the entire 671–319 μm range of the SLW detector.

respectively, although the $^{13}\text{CO } T_{\text{rot}}$ is intermediate between these two. Simple LVG modeling shows the presence of warm (125 K) and dense ($n_{\text{H}_2} \sim 7 \times 10^4 \text{ cm}^{-3}$) gas, which is traced by in the SW extension in the CO and CI maps. The observed flux from the 398.6 μm H_2O line is consistent with these values.

Acknowledgements. SPIRE has been developed by a consortium of institutes led by Cardiff Univ. (UK) and including Univ. Lethbridge (Canada); NAOC (China); CEA, LAM (France); IFSI, Univ. Padua (Italy); IAC (Spain); Stockholm Observatory (Sweden); Imperial College London, RAL, UCL-MSSL, UKATC, Univ. Sussex (UK); Caltech, JPL, NHSC, Univ. Colorado (USA). This has been supported by national funding agencies: CSA (Canada); NAOC (China); CEA, CNES, CNRS (France); ASI (Italy); MCINN (Spain); Stockholm Observatory (Sweden); STFC (UK); and NASA (USA).

References

- Buckle, J. V., Curtis, E. I., Roberts, J. F., et al. 2010, MNRAS, 401, 204
- Cruz-González, I., Salas, L., & Hiriart, D. 2007, Rev. Mex. Astron. Astrofis, 43, 337
- Fixen, D. J., Bennett C. L., & Mather, J. C. 1999, ApJ, 526, 207
- Garden, R., Hayashi, M., Gatley, I., et al. 1991, ApJ, 374, 540
- Griffin, M. J., et al. 2010, A&A, 518, L3
- Hjalmarson, A., Frisk, U., Olberg, M., et al. 2003, A&A, 402, L39
- Jakob, H., Kramer, C., Simon, R. et al., 2007, A&A, 461, 999
- Jones, S. C., Naylor, D. A., Gom, B. G., et al. 2009, Proc. 30th Canadian Symposium on Remote Sensing, Canadian Remote Sensing Society, in press
- Liechti, S., & Walmsley, C. M. 1997, A&A, 321, 625
- Oberst, T. E., Parshley, S. C., Stacey, G. J., et al. 2006, ApJ, 652, 125
- Pilbratt, G. L., et al. 2010, A&A, 518, L1
- Richardson, K. J., White, G. J., et al. 1986, MNRAS, 219, 167, 1986
- Richardson, K. J., White, G. J., Monteiro, T., et al. 1988, A&A, 198, 237
- Richardson, K. J., Sandell, G., & Krisciunas, K. 1989, A&A, 224, 199
- Schneider, N., Bontemps, S., Simon, R., et al. 2006, A&A, 458, 855
- Schneider, N., Csengeri, N., Bontemps, S., et al. 2010, A&A, accepted [arXiv:1003.4198]
- Swinyard, B. M., et al. 2010, A&A, 518, L4
- Van der Tak, F., Black, J., Schoier, F., et al. 2007, A&A, 468, 627
- Wilson, T. L., & Mauersberger, R. 1990, A&A, 239, 305

-
- ¹ Department of Physics & Astronomy, The Open University, UK
e-mail: g.j.white@open.ac.uk
- ² Space Science Department, Rutherford Appleton Laboratory, Chilton, UK
- ³ Institute for Space Imaging Science, University of Lethbridge, Lethbridge, Alberta, Canada
- ⁴ Institut d'Astrophysique Spatiale, CNRS/Université Paris-Sud 11, 91405 Orsay, France
- ⁵ Laboratoire AIM, CEA/IRFU – CNRS/INSU – Université Paris Diderot, CEA-Saclay, 91191 Gif-sur-Yvette Cedex, France
- ⁶ CEA, Laboratoire AIM, Irfu/SAP, Orme des Merisiers, 91191 Gif-sur-Yvette, France
- ⁷ Department of Physics and Astronomy, Cardiff University, Cardiff, UK
- ⁸ Laboratoire d'Astrophysique de Marseille, UMR6110 CNRS, 38 rue F. Joliot-Curie, 13388 Marseille, France

- ⁹ Canadian Institute for Theoretical Astrophysics, Toronto, Ontario, M5S 3H8, Canada
- ¹⁰ University of California, Radio Astronomy Laboratory, Berkeley, 601 Campbell Hall, US Berkeley CA 94720-3411, USA
- ¹¹ Institut de Radioastronomie Millimétrique (IRAM), 300 rue de la Piscine, 38406 Saint-Martin-d'Hères, France
- ¹² Joint Astronomy Centre, University Park, Hilo, USA
- ¹³ National Astronomical Observatories (China)
- ¹⁴ Istituto di Fisica dello Spazio Interplanetario, INAF, Via del Fosso del Cavaliere 100, 00133 Roma, Italy
- ¹⁵ NASA-Goddard SFC, USA
- ¹⁶ Laboratoire des Signaux et Systèmes (CNRS Supélec Université Paris-Sud 11), Plateau de Moulon, 91192 Gif-sur-Yvette, France
- ¹⁷ Université de Toulouse; UPS; CESR; 9 avenue du colonel Roche, 31028 Toulouse Cedex 4, France
- ¹⁸ CNRS; UMR5187; 31028 Toulouse, France